

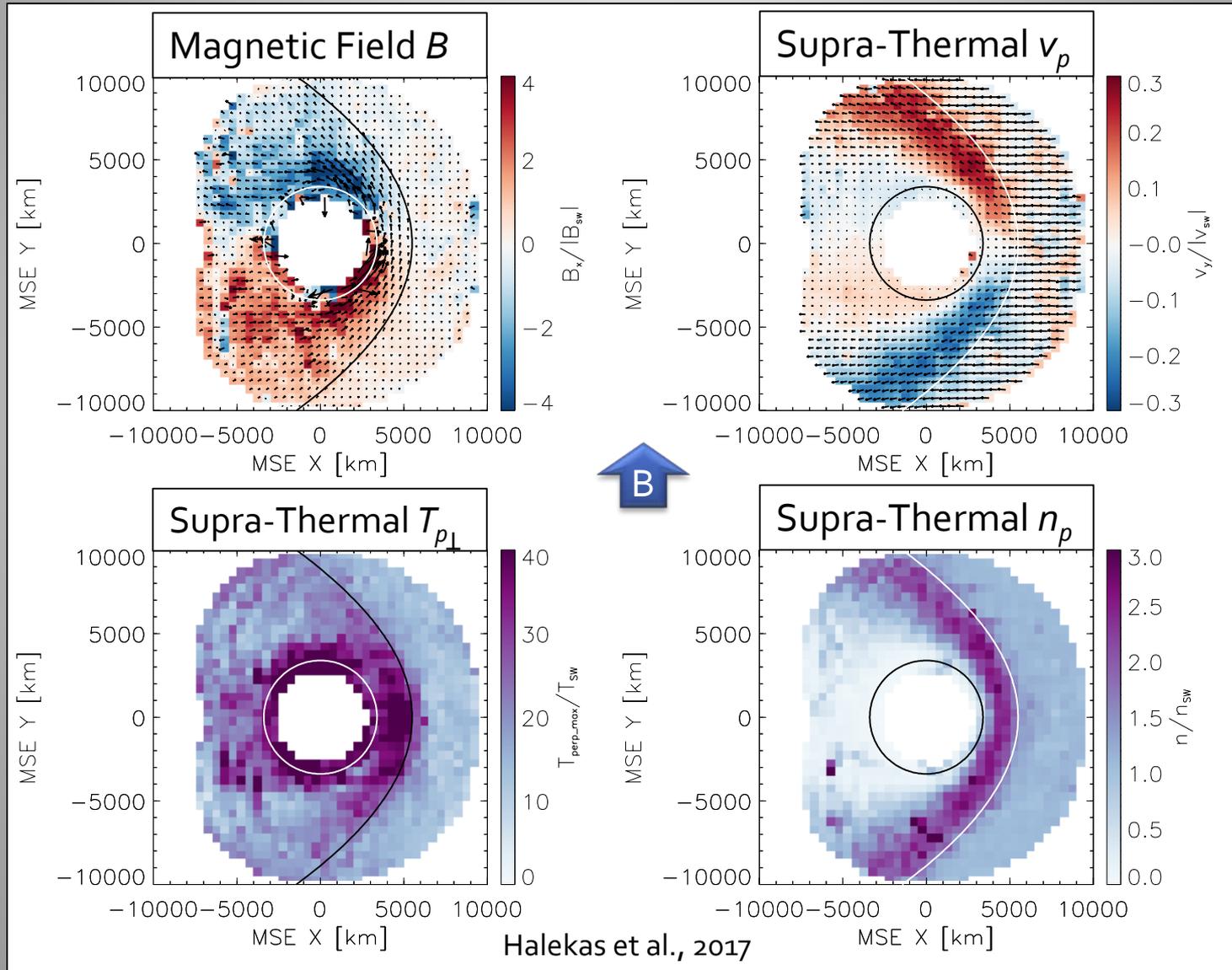


Momentum Transfer and Boundary Layer Structure at Mars

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Mars-Solar Wind Interaction



Multi-Fluid Force Model

Momentum Coupling
(Pickup, Mass-Loading)

Magnetic
Pressure/
Tension

Particle
Pressure

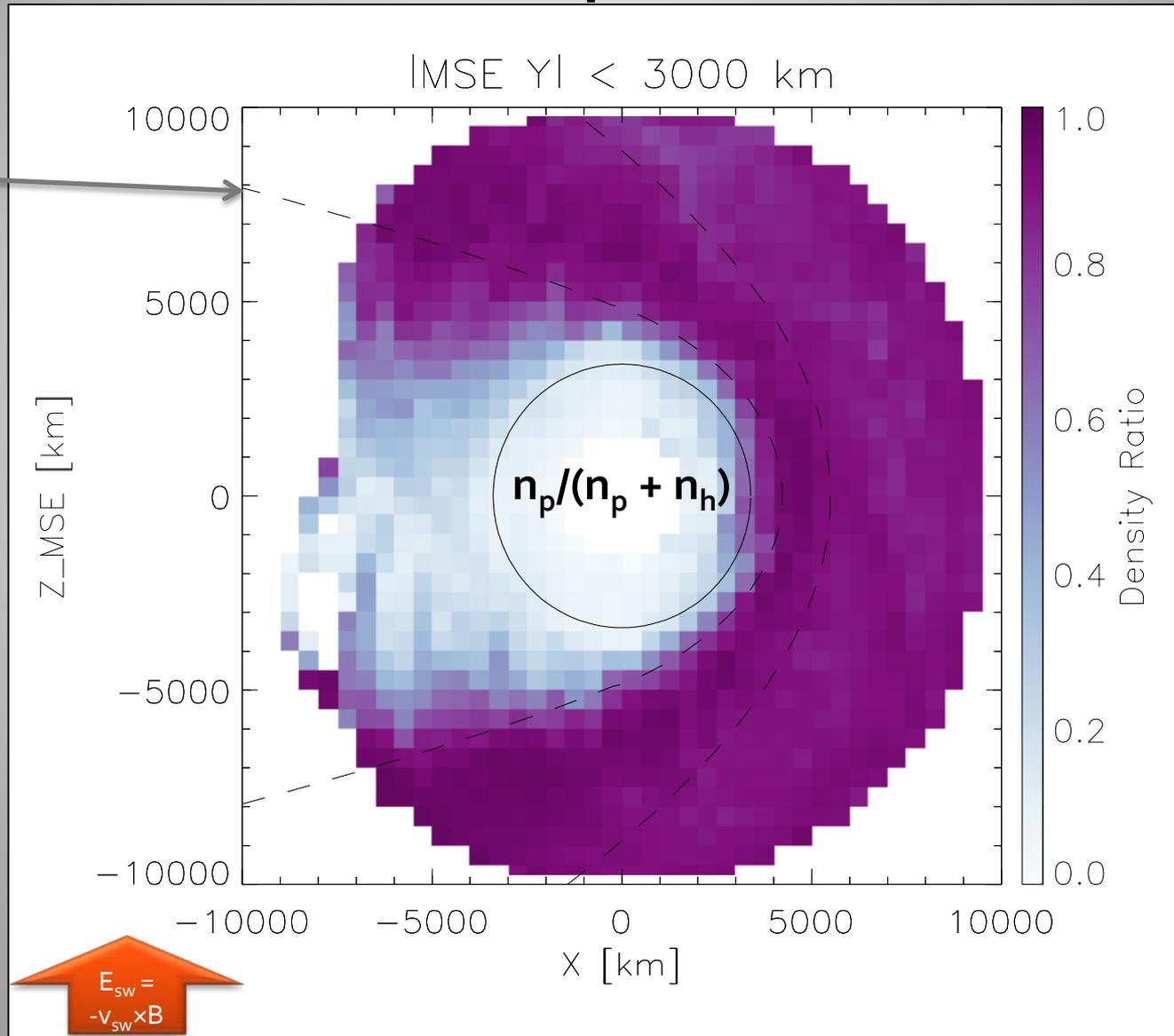
$$\vec{F}_p = m_p n_p \left(\frac{\partial}{\partial t} + \vec{v}_p \cdot \nabla \right) \vec{v}_p = \frac{n_p}{n_e} \left[q n_h (\vec{v}_p - \vec{v}_h) \times \vec{B} + \vec{J} \times \vec{B} - \nabla P_e \right] - \nabla \cdot \vec{P}_p$$

$$\vec{F}_h = m_h n_h \left(\frac{\partial}{\partial t} + \vec{v}_h \cdot \nabla \right) \vec{v}_h = \frac{n_h}{n_e} \left[q n_p (\vec{v}_h - \vec{v}_p) \times \vec{B} + \vec{J} \times \vec{B} - \nabla P_e \right] - \nabla \cdot \vec{P}_h$$

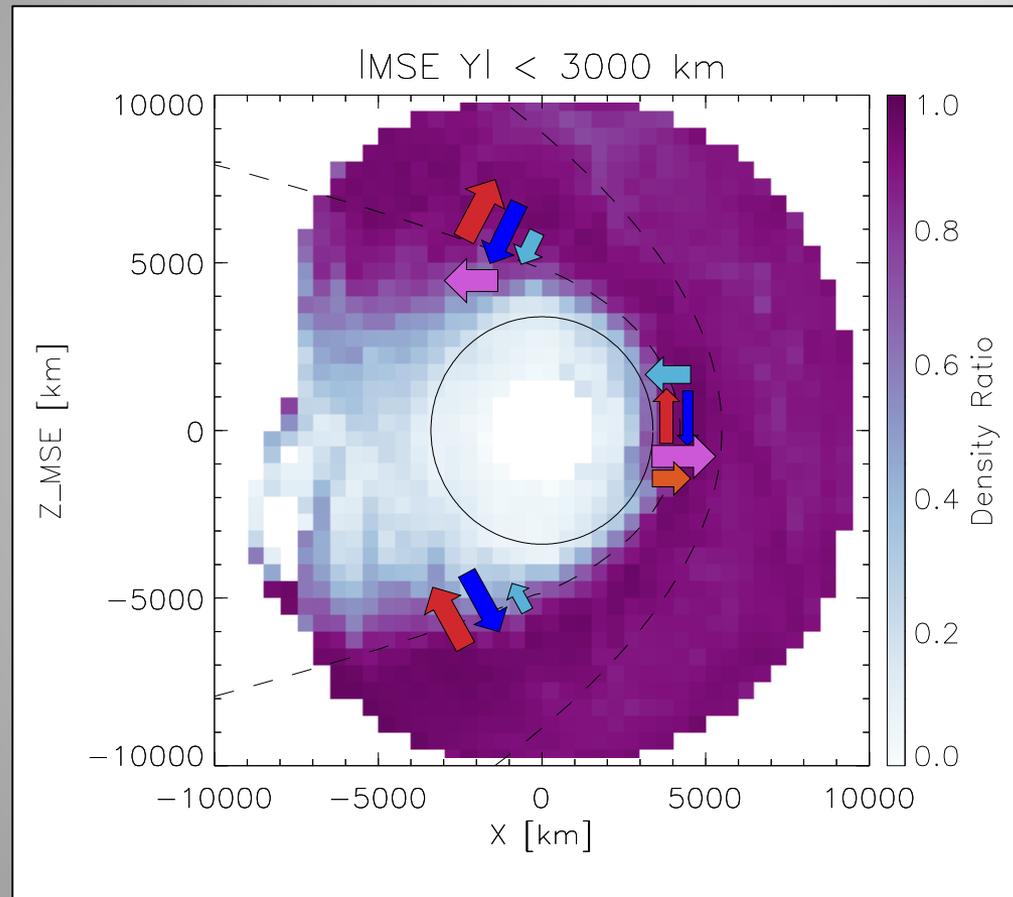
First order model for solar wind protons interacting with heavy planetary ions

Ion Composition

Nominal MPB
[Trotignon
et al., 2006]



Ion Composition Boundary Forces



Thicker
boundary
expected
in $+E_{sw}$

Narrower
boundary
expected
in $-E_{sw}$

$\mathbf{v} \times \mathbf{B}$ force on heavies

$\text{Grad}P_h$ force on heavies

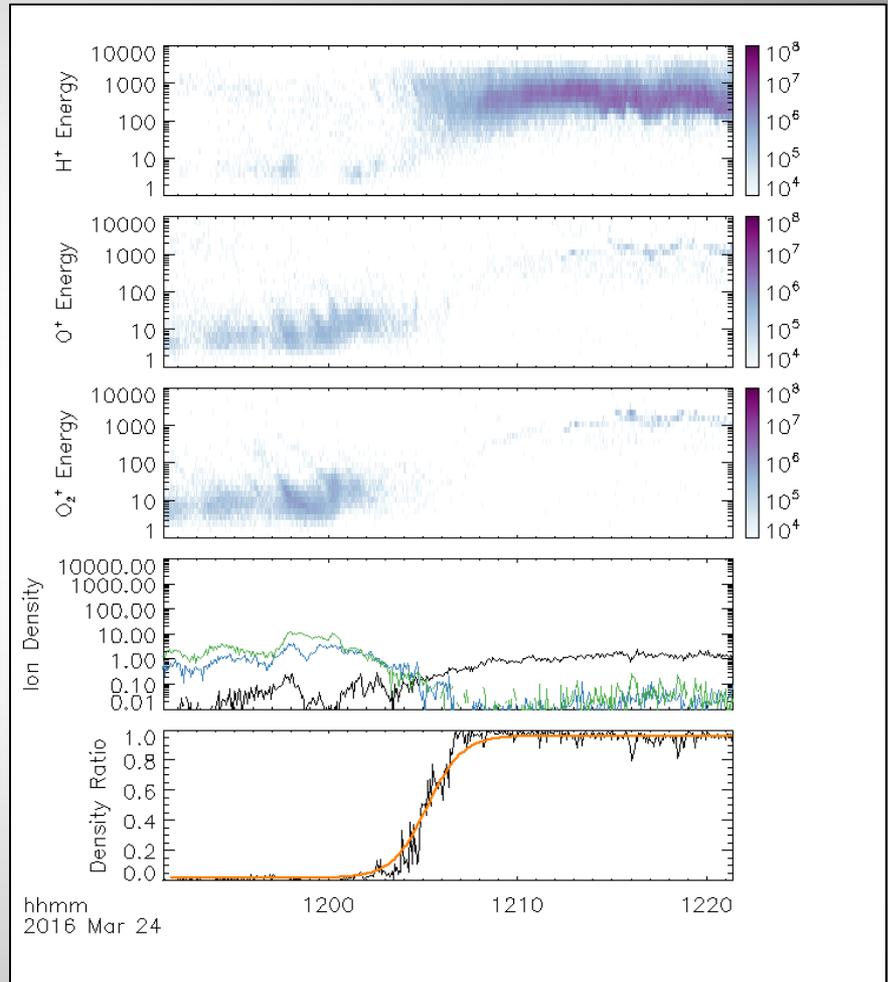
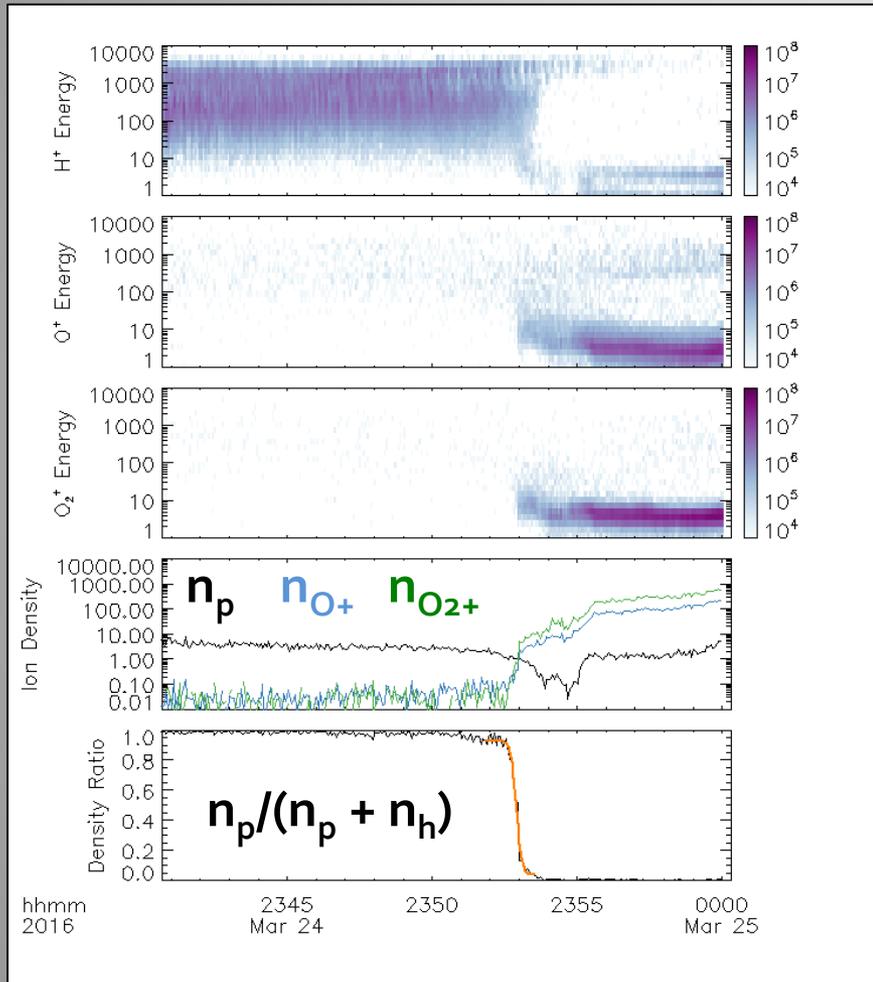
$\mathbf{v} \times \mathbf{B}$ force on protons

$\text{Grad}P_p$ force on protons

$\mathbf{J} \times \mathbf{B}$ force on all ions

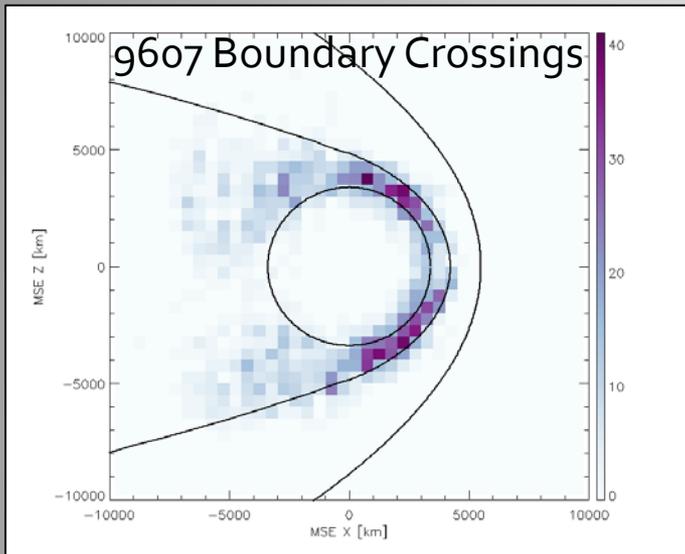
$\text{Grad}P_e$ force on all ions?

Boundary Layer Fits

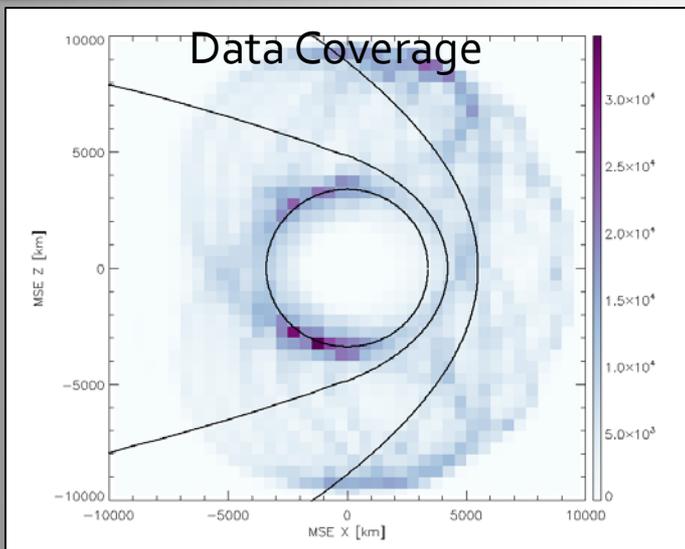
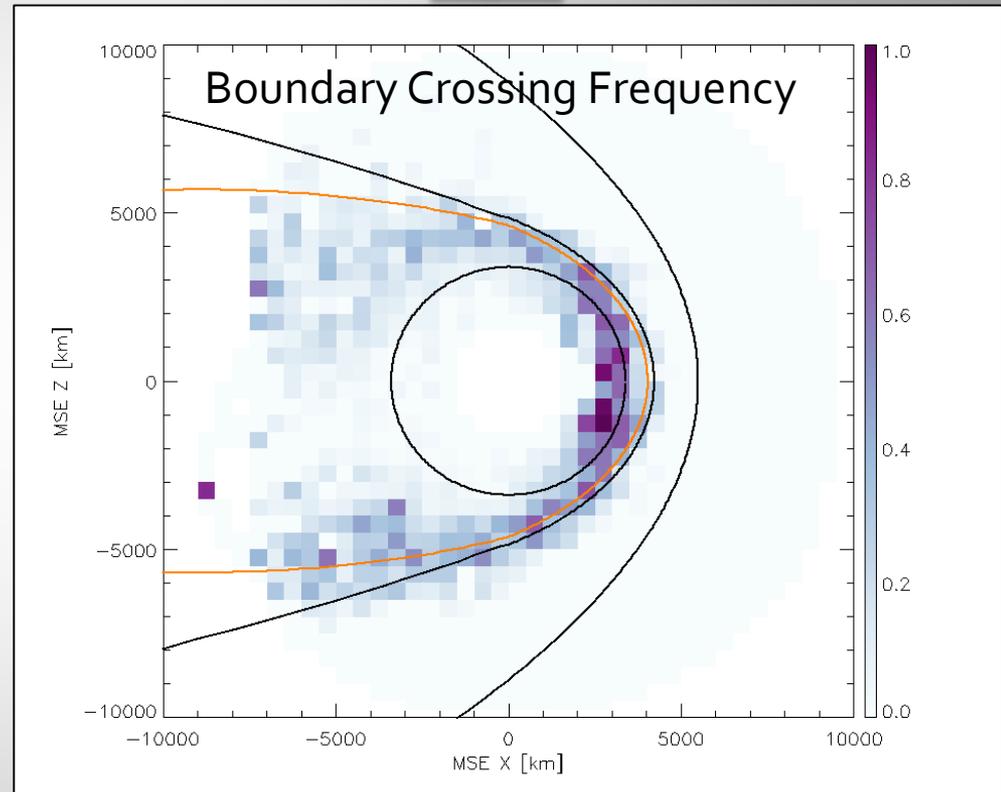


- Fit boundary density ratio to functional form $f(t) = A_0 + A_1 * \tanh[(t - A_2) / A_3]$
- Convert $2 * A_3$ to spatial thickness using spacecraft velocity normal to average boundary

Boundary Layer Position

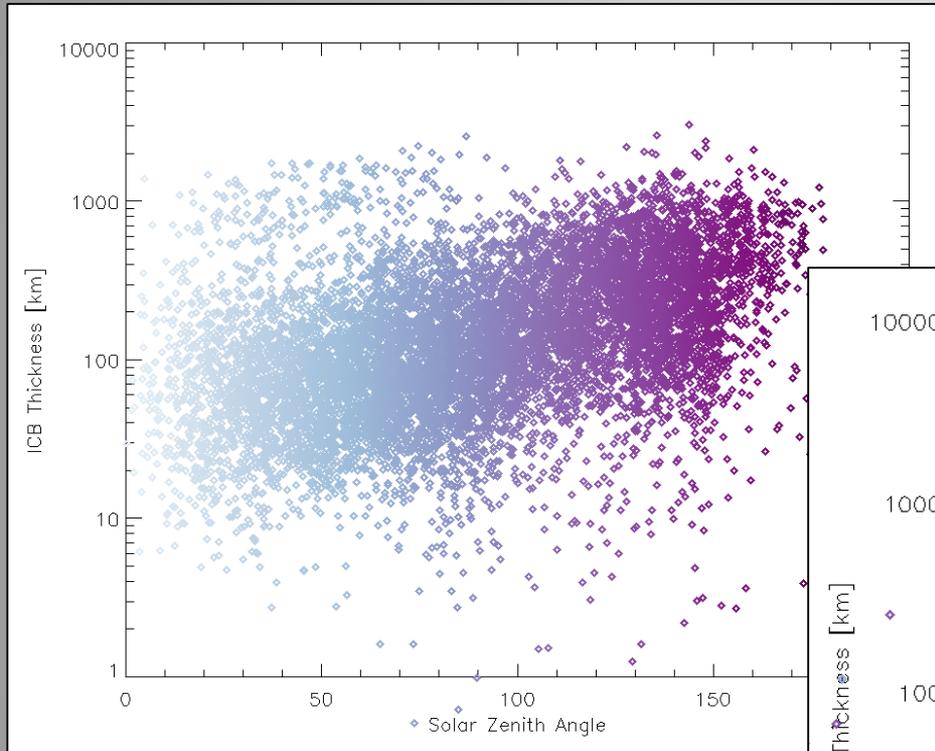


$$E_{sw} = -v_{sw} \times B$$

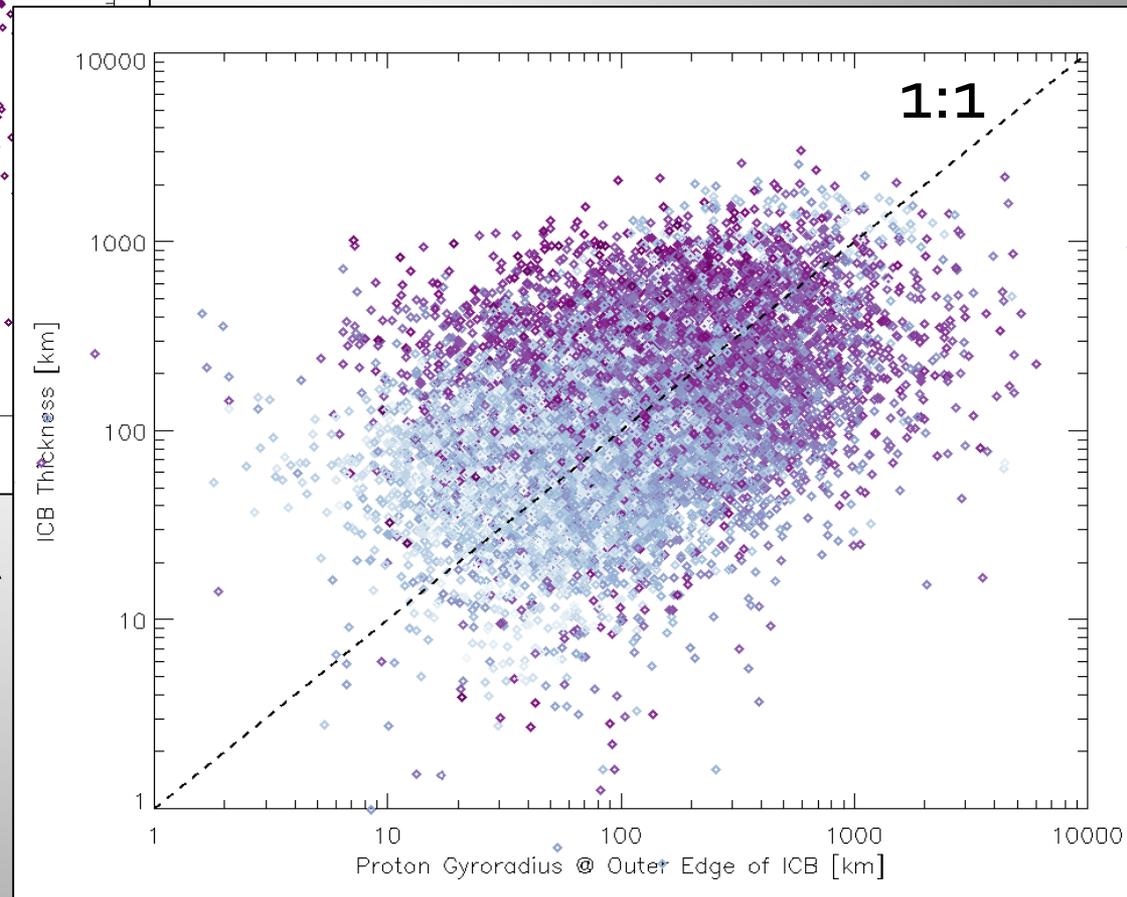


- Average ICB close to nominal MPB, but less flared in tail
- ICB asymmetric in MSE

Boundary Layer Thickness



Thickness scales with gyroradius

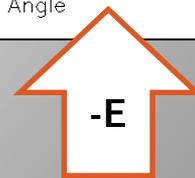
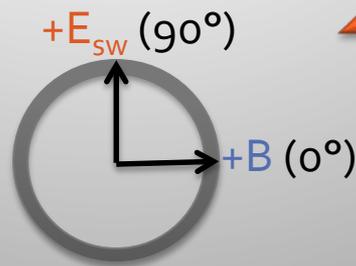
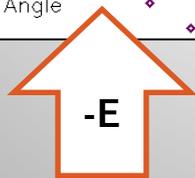
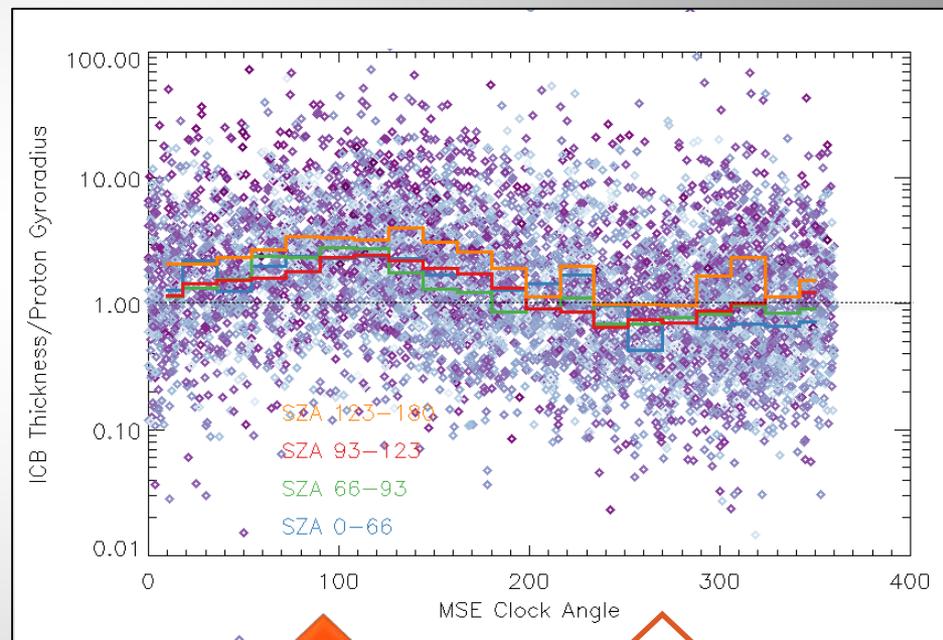
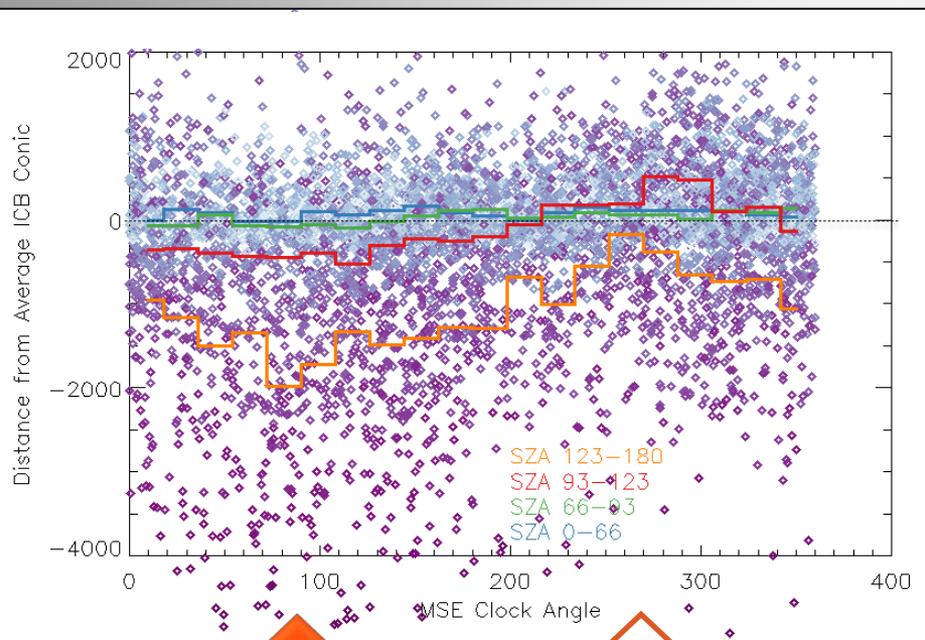


Thickness increases with SZA

Boundary Layer Asymmetries

Boundary farther out in -E hemisphere

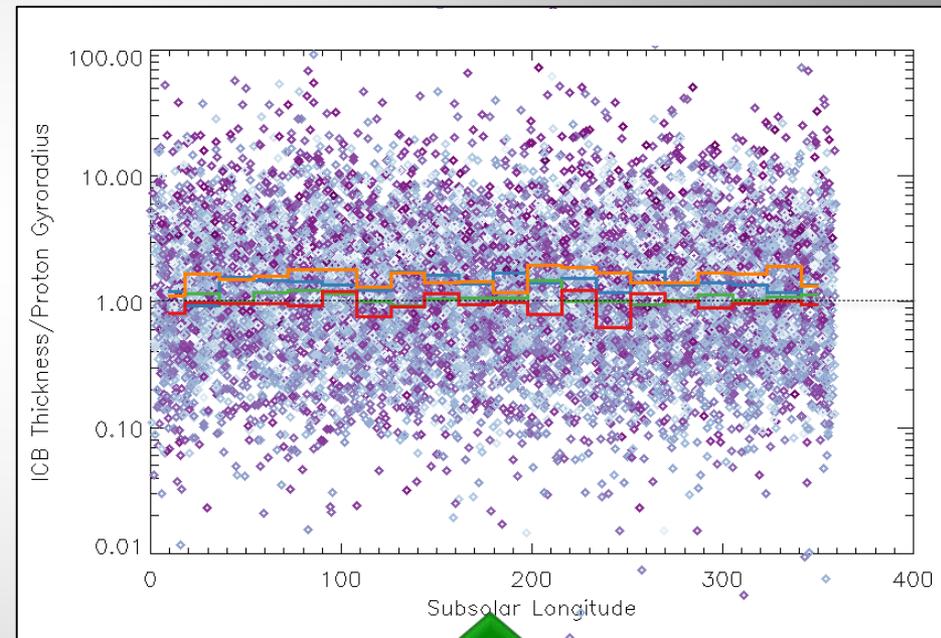
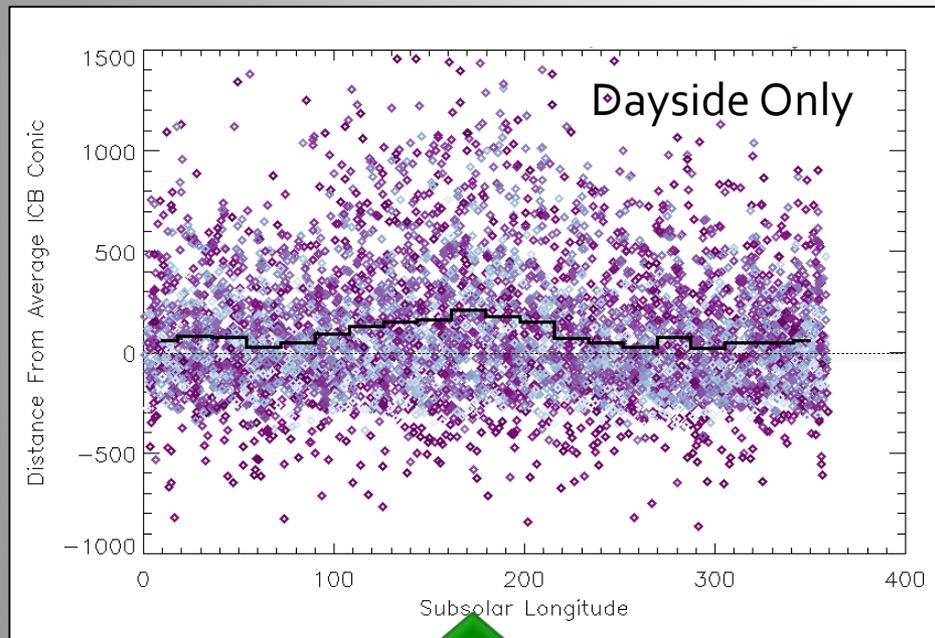
Boundary thicker in +E hemisphere



Boundary Layer Asymmetries II

Boundary farther out over
crustal magnetic fields

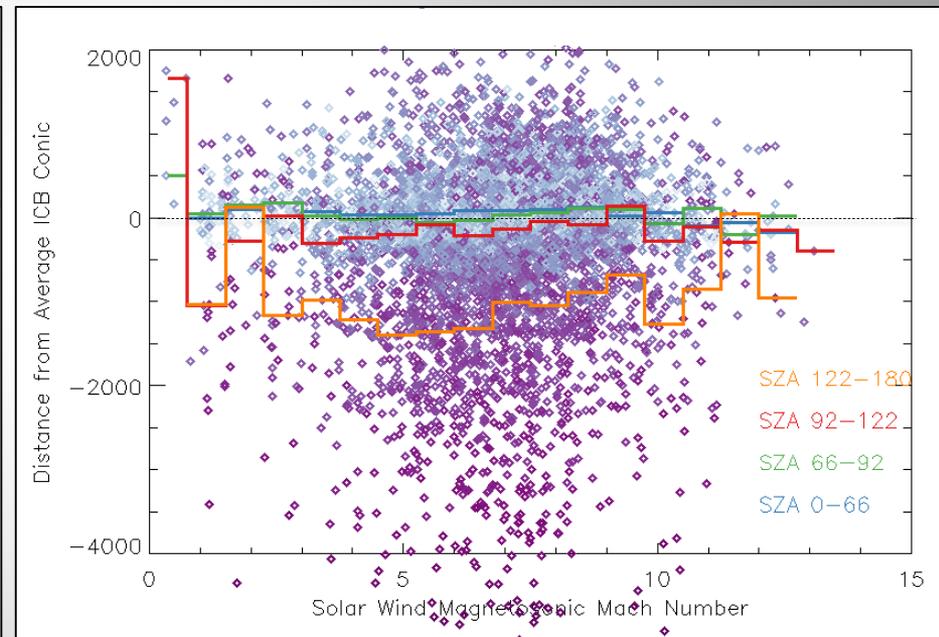
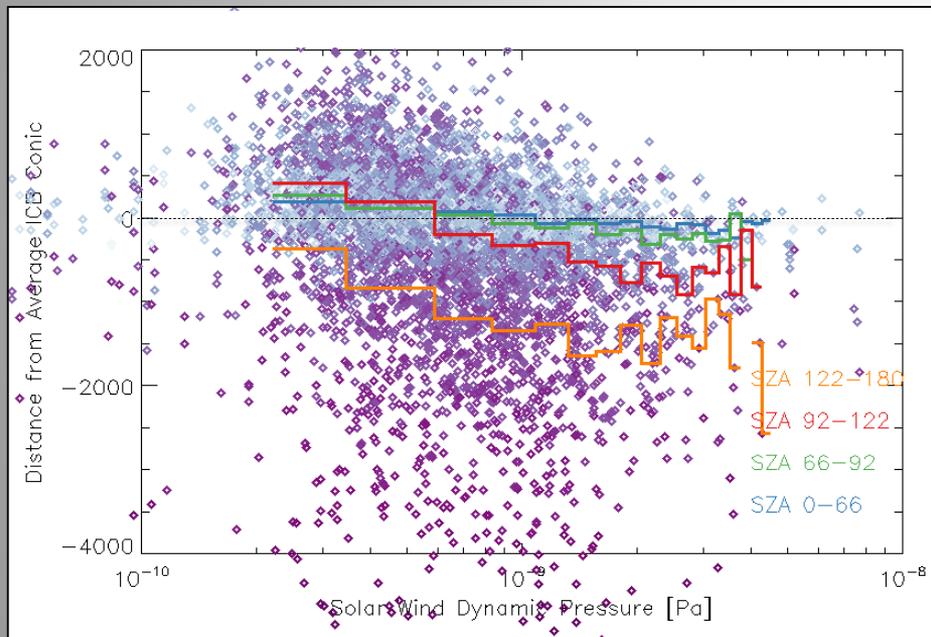
Boundary thickness unaffected
by crustal magnetic fields



Boundary Layer Solar Wind Dependence

Boundary compressed by solar wind ram pressure

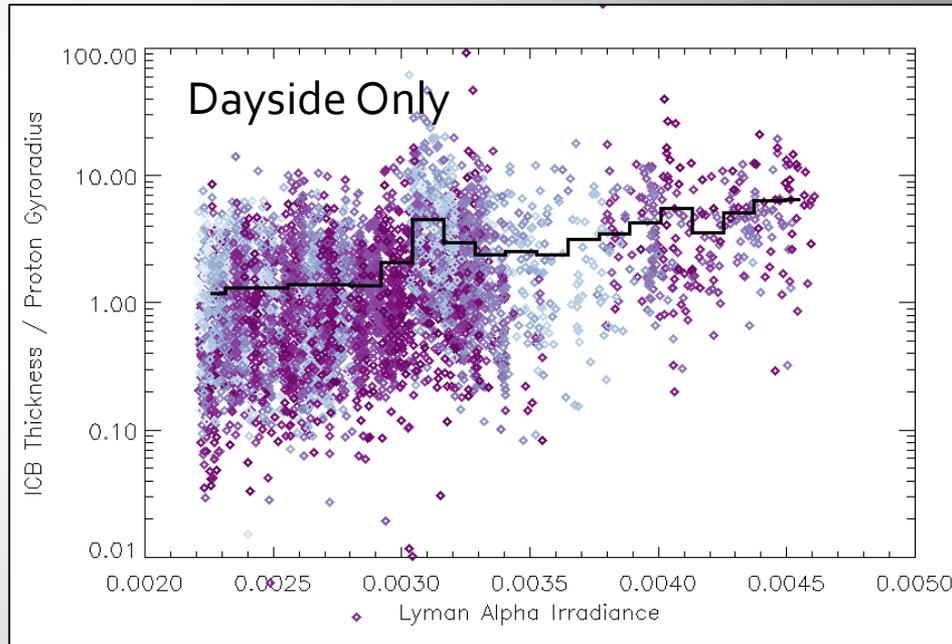
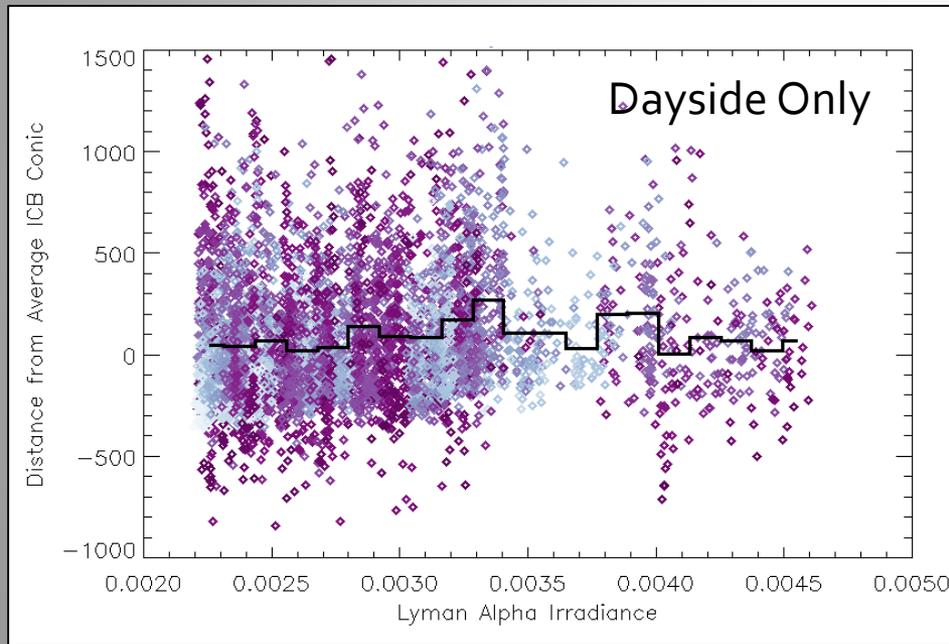
Boundary insensitive to solar wind Mach number



Boundary Layer EUV Dependence

Boundary position doesn't
depend on EUV...

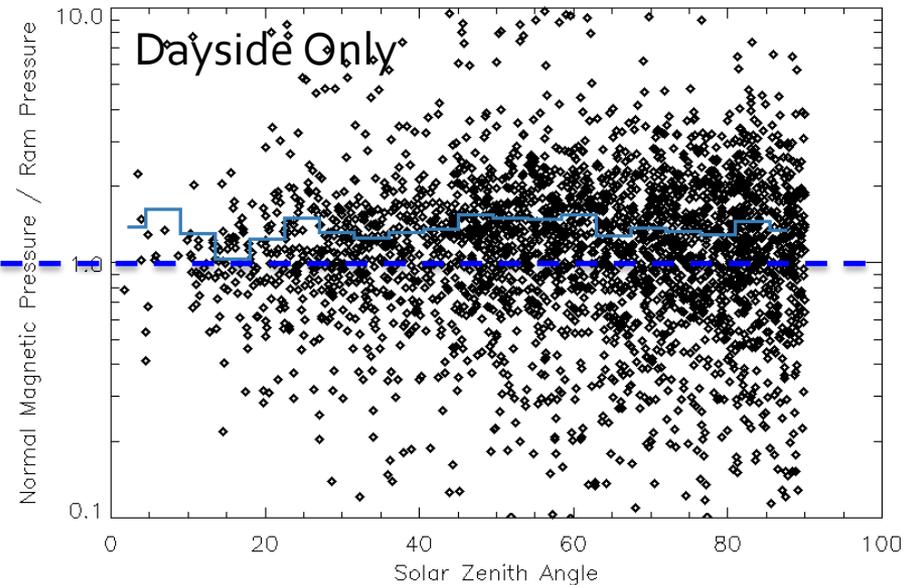
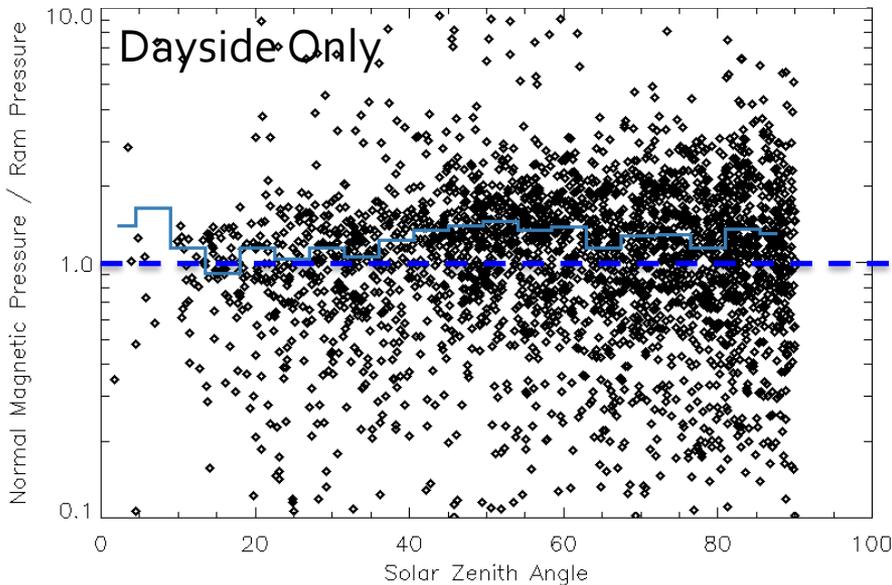
But boundary thickness
does!



Pressure Balance at Boundary Layer

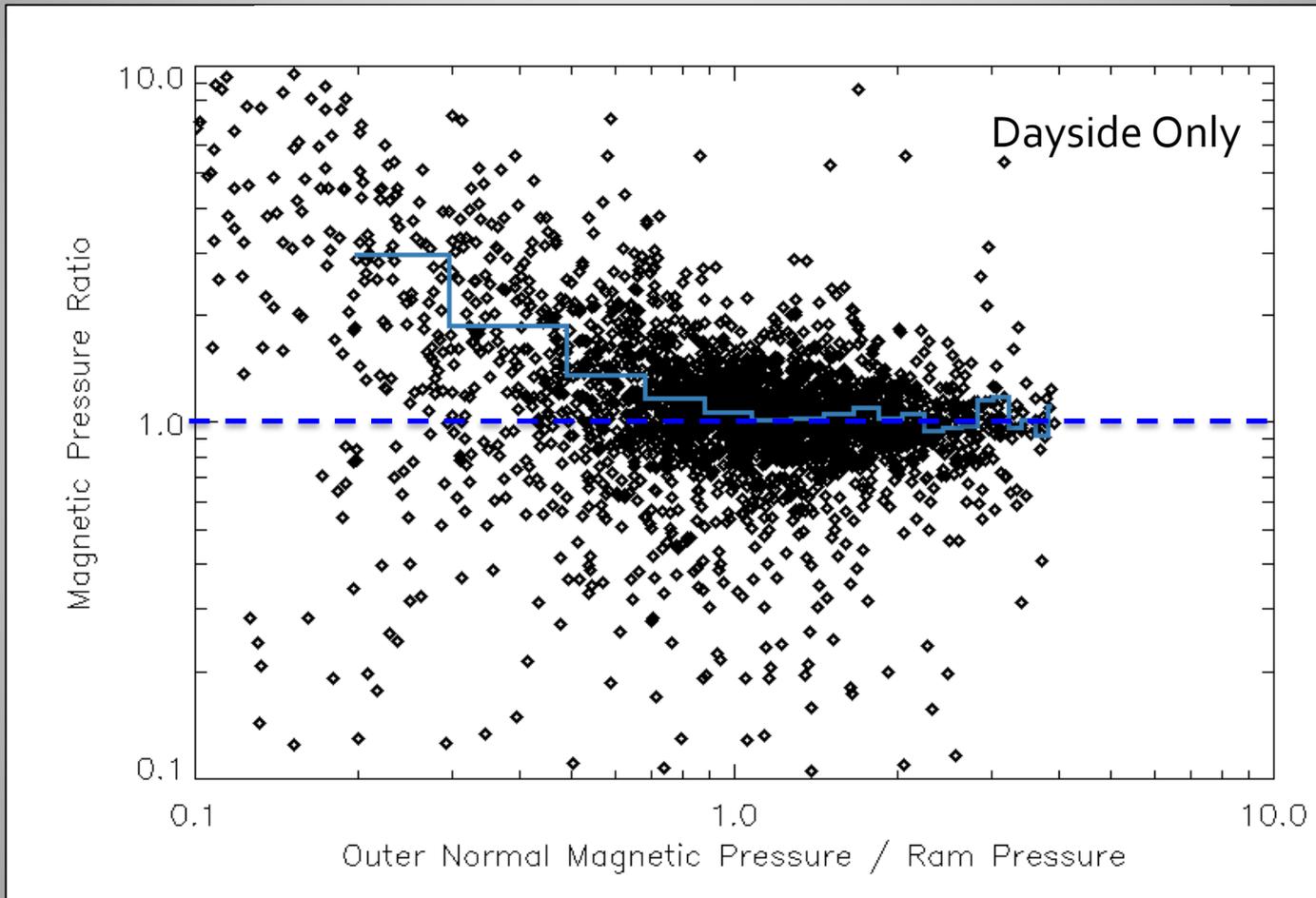
Outer ion composition boundary layer

Inner ion composition boundary layer



- Tangential magnetic field balances normal component of solar wind ram pressure
- Pressure balance already established by the outer edge of the composition boundary
- Excess of magnetic pressure – result of crustal magnetic fields?

Magnetic Field Jump

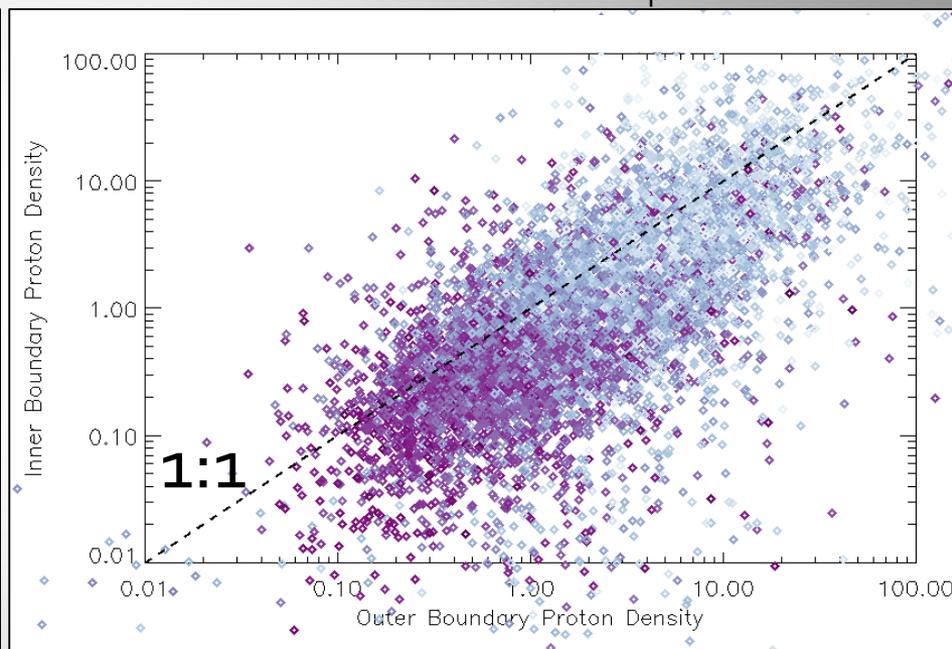
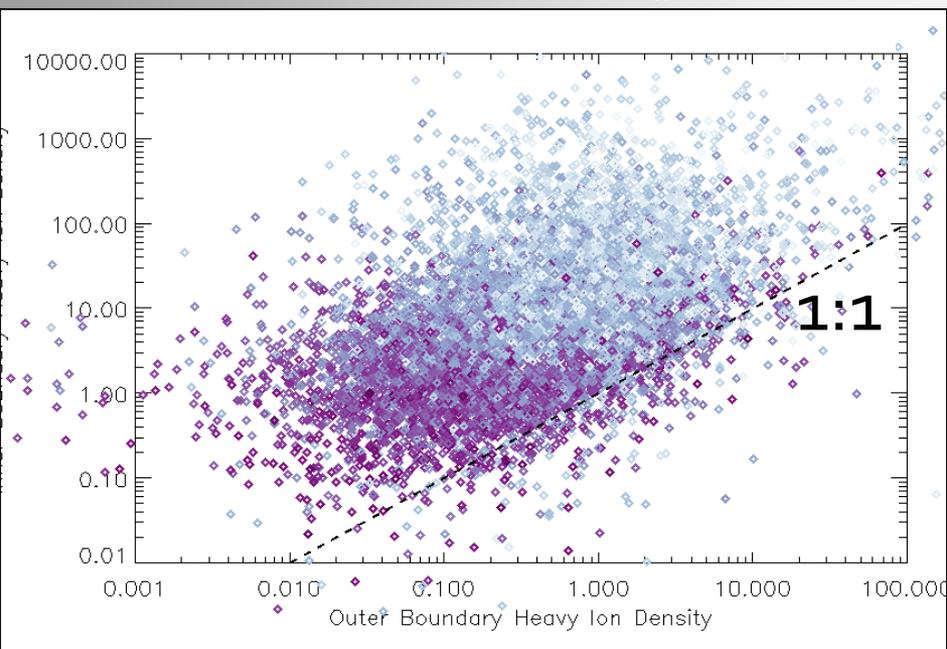


- Magnetic field and pressure jump across boundary usually moderate
- ICB \neq MPB (at least not the same thickness!)

Density Jump

Heavy Ion Density n_h

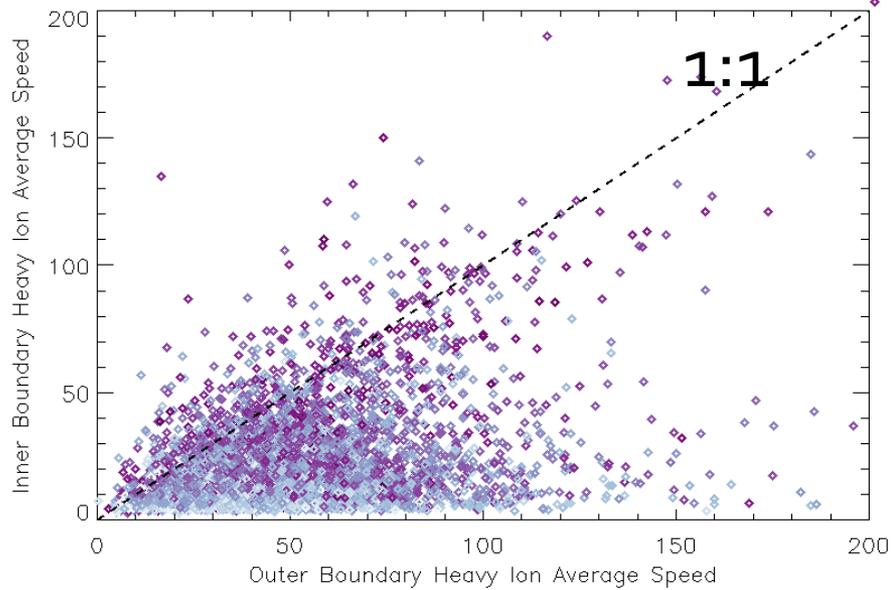
Proton Density n_p



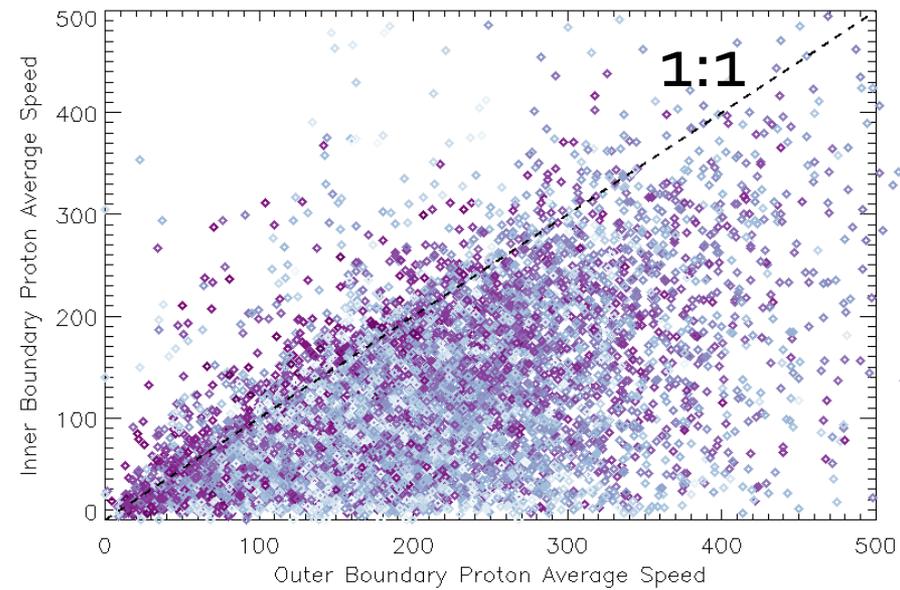
- Heavy ion density increases across boundary
- Proton density usually drops across boundary

Speed Jump

Heavy Ion Speed $\langle v_h \rangle$



Proton Speed $\langle v_p \rangle$



- Both heavy ion and proton speed are lower inside the boundary

Conclusions

- The transition in composition between solar wind and planetary ions is part of a complex boundary layer
 - The ion composition boundary (ICB) is not (always) the same as either the MPB or the ionopause
 - The ICB has significant asymmetries in both location and thickness, ultimately resulting from the asymmetry of the Lorentz force