

Eclipsing the X-ray emitting regions in AGN

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In collaboration with:

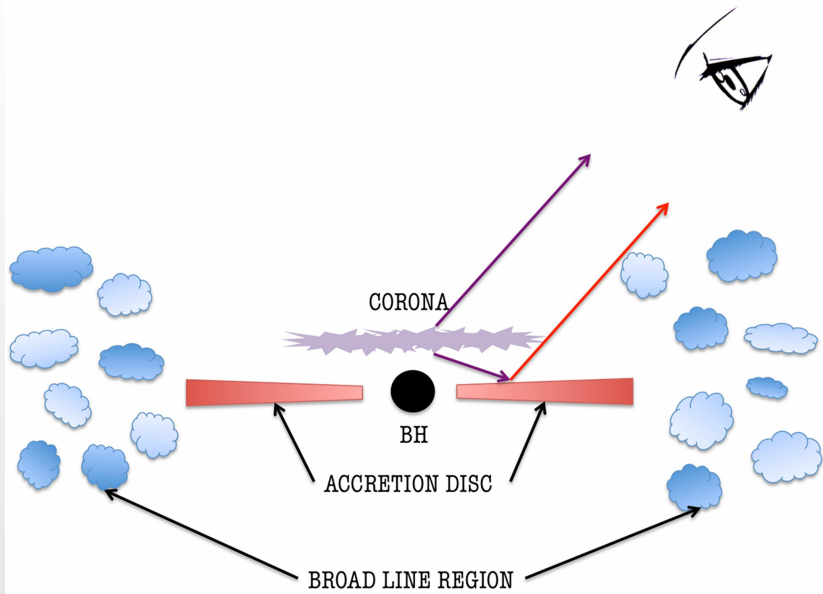
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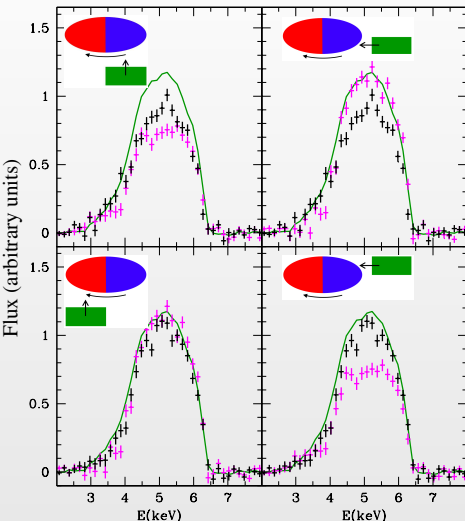
The Extremes of Black Hole Accretion | ESAC, June 8, 2015

The innermost regions in AGN: general scheme



Motivation

To probe general relativistic effects by studying X-ray eclipses



NGC1365

Risaliti et al. +11

(MNRAS 417, 178–183)

Compton–thin cloud

$$N_{\text{H}} = 3.5 \times 10^{23} \text{cm}^{-2}$$

$$A_{\text{eff}} = 2 \text{m}^2 @ 6 \text{keV}$$

The KYNCONV model: intro

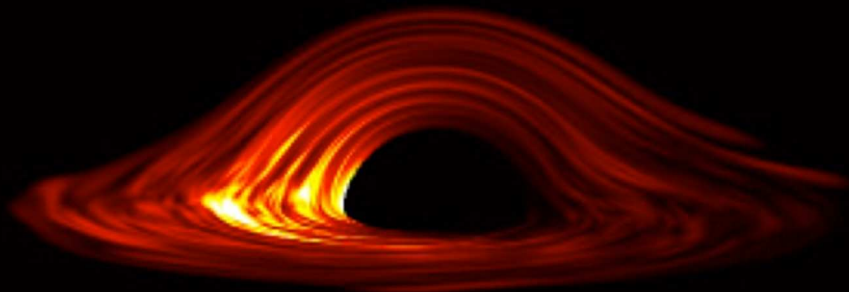
Set of κY relativistic models in Dovčiak et al. +04 (ApJS 153, 205–221)

relativistic convolution model

accretion discs spectra

strong gravity regime

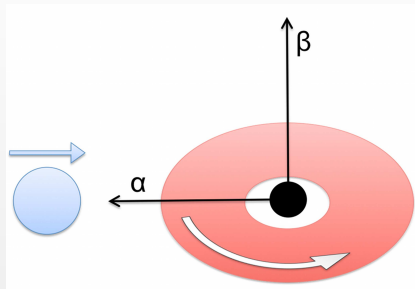
obscuration with a circular cloud



The KYNCONV model: main parameters

System inclination $\rightarrow \theta_0$ /deg

Central engine $\left\{ \begin{array}{ll} \text{BH spin} & \rightarrow a \\ \text{emissivity index} & \rightarrow q \end{array} \right.$

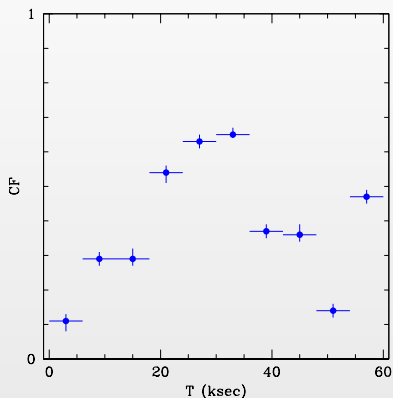


Cloud $\left\{ \begin{array}{ll} \text{position} & \rightarrow \alpha, \beta / \frac{GM}{c^2} \\ \text{size} & \rightarrow r / \frac{GM}{c^2} \\ \text{column density} & \rightarrow N_H / \text{cm}^{-2} \text{ (phabs, zxipcf)} \\ \text{ionisation} & \rightarrow \log \xi, \text{ with } \xi / \frac{\text{erg cm}}{\text{s}} \text{ (zxipcf)} \end{array} \right.$

What we could expect

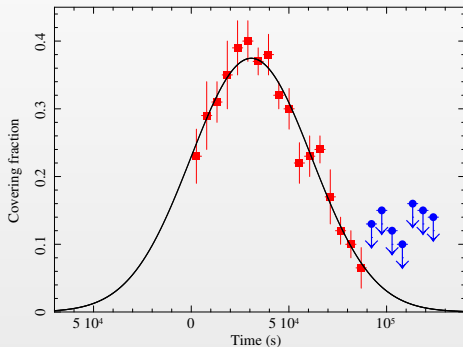
Previous detections of disc-emission anisotropies?

NGC1365



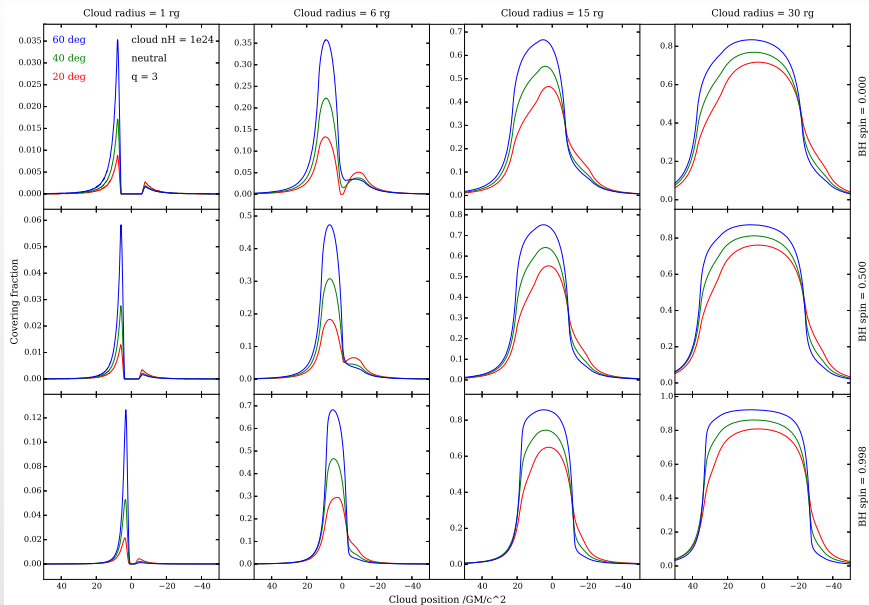
Risaliti et al. +09 (ApJ 696, 160–171)

SWIFT J2127.4+5654

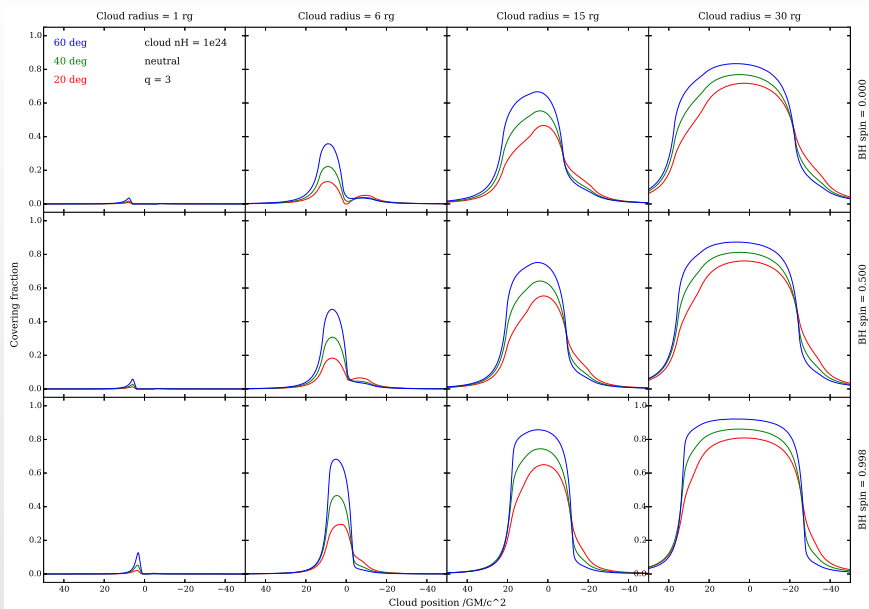


Sanfrutos et al. +13 (MNRAS 436, 1588–1594)

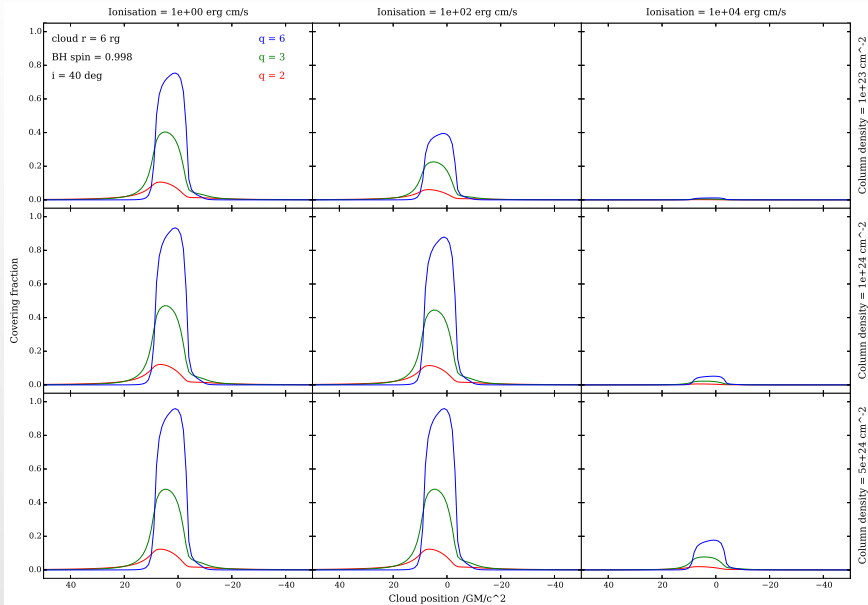
The covering fraction profiles (I)



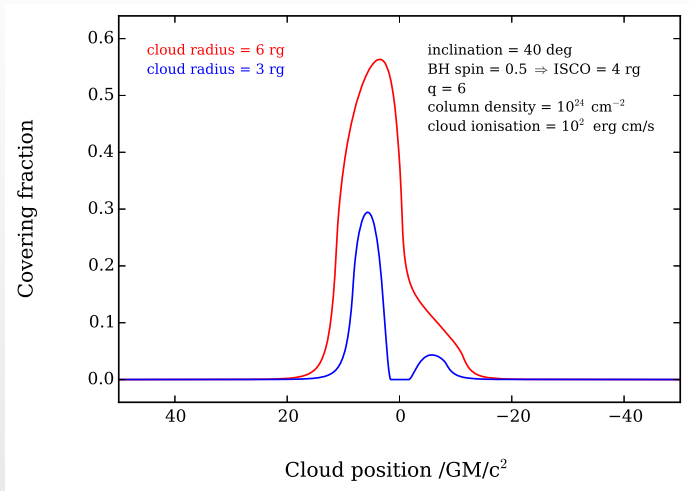
The covering fraction profiles (I)



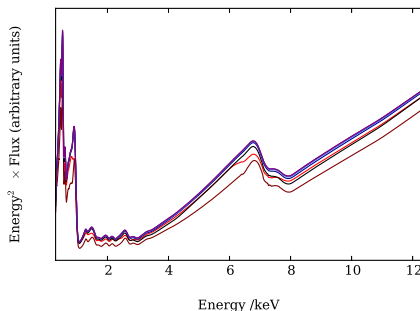
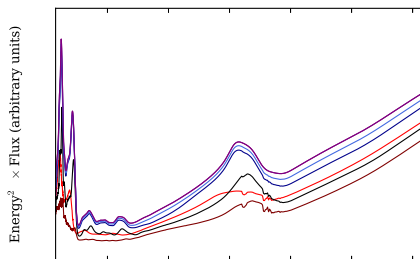
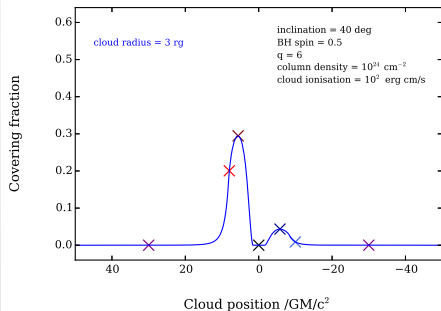
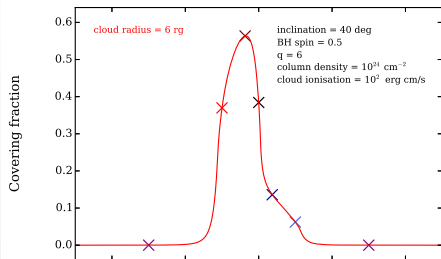
The covering fraction profiles (II)



The covering fraction profiles (III)



Detectability on spectral features



Conclusions

- Preliminary results are promising
- Tentative detections with current instruments (NGC1365, SWIFT J2127.4+5654)
- Ongoing and future work
 - Fit real covering fraction profiles
 - Fit real spectra
 - Simulate future instruments' spectra (ASTRO-H, ATHENA+)
 - Check degeneracies
 - ...

... SO MUCH WORK TO DO!