

THE EARTH AS AN EXOPLANET



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Infrared terrestrial emission

$$F_{abs} = F_{emit}$$

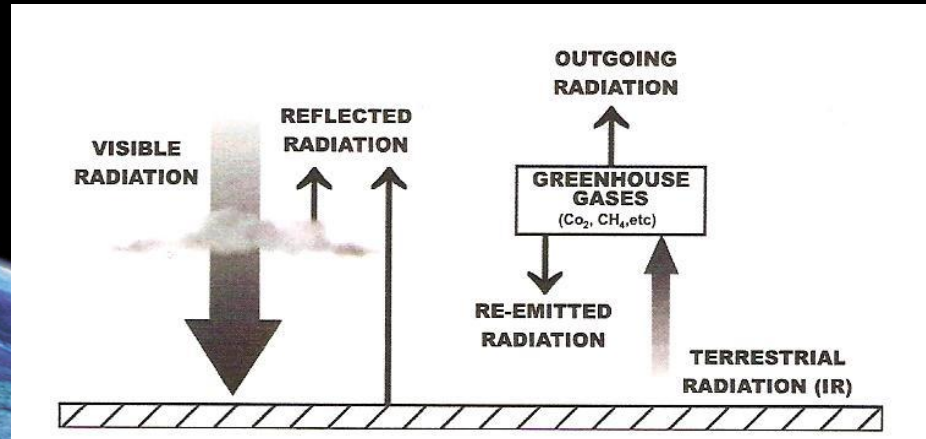
$$(\pi R_E^2) F_S (1 - a) = 4\pi R_E^2 F_E$$

$$F_E = (1 - g) \sigma T^4$$

$$T_{eq}^4 = \frac{F_S (1 - a)}{4\sigma(1 - g)} = \frac{L_S (1 - a)}{4d^2\sigma(1 - g)}$$

$$T \approx 288 \text{ K}$$

$$\lambda = \frac{2898 \mu\text{m}}{288 \text{ K}} \approx 10 \mu\text{m}$$



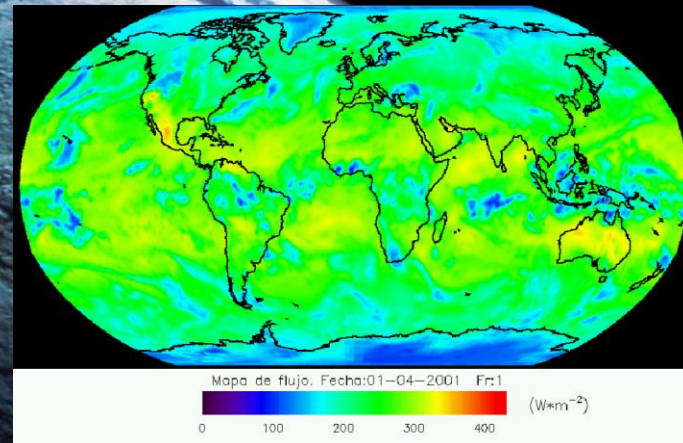
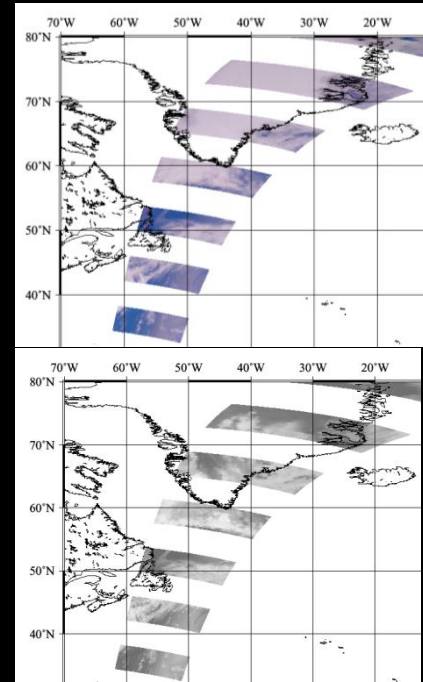
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* **GEWEX** Global Energy and Water-cycle Experiment (NASA)

SRB Surface Radiation Budget

LW Longwave radiation ($\lambda > 4.0 \mu\text{m}$) $\epsilon = \pm 5 \text{ W/ m}^2$

- * Data period: 1986 - 2001
- * Observations each 3 hours:
(0, 3, 6, 9, 12, 15, 18, ,21) hrs UT per day.
- * Numerical matrices of $1^\circ \times 1^\circ$ (lon,lat) resolution.
transformed in emission flux maps.

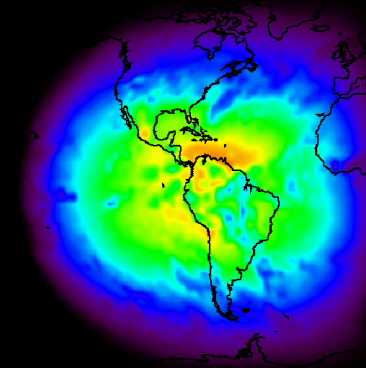


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Geometrical model

Planetary disk simulation

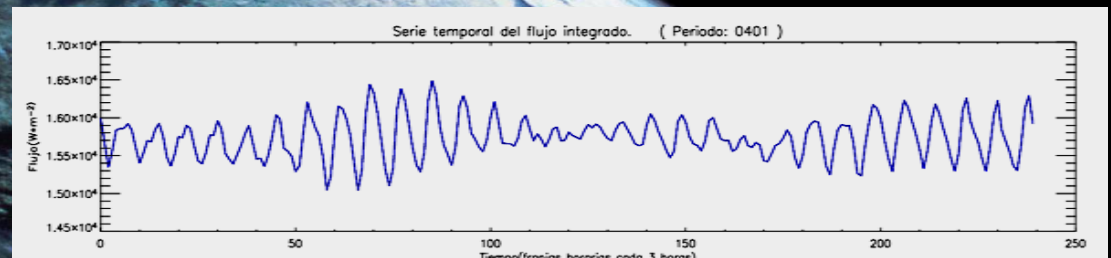
Weight map= flux map*area map*cos map



Point-like signal

Integrated flux= \int (Weight map) ds

It's transformed in a flux vector
-> time series

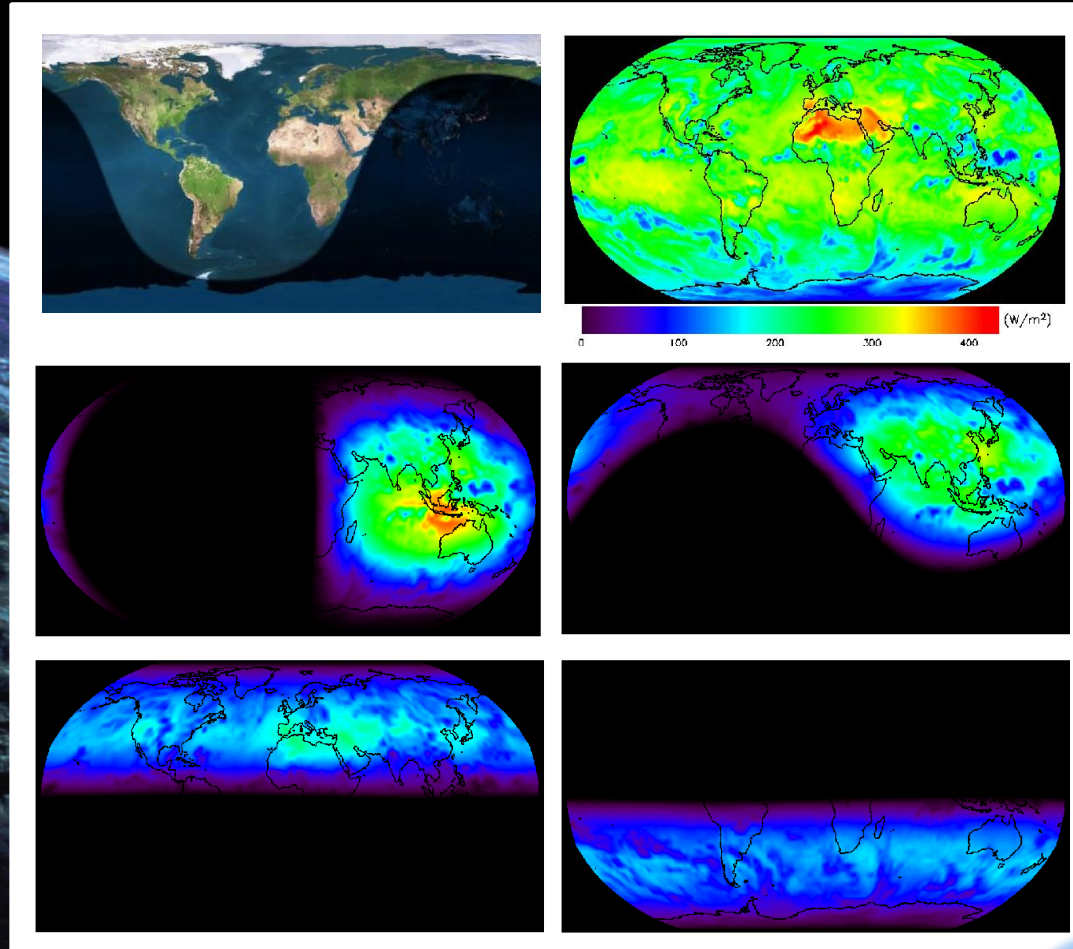


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Variation of the observer's geometry

Inclination angles

Equator, 45° N, 90° N, 90° S.



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Time series

Annual maximum (~5%) in boreal summer.

“Northern-hemisphere” behaviour.

Fig. 5.— Time series. From top to bottom, the years of 1987, 2001 and January, April, July and October 2001 for 90°, 45°, 0° and -90° inclination angles (black, red, blue and green lines respectively).

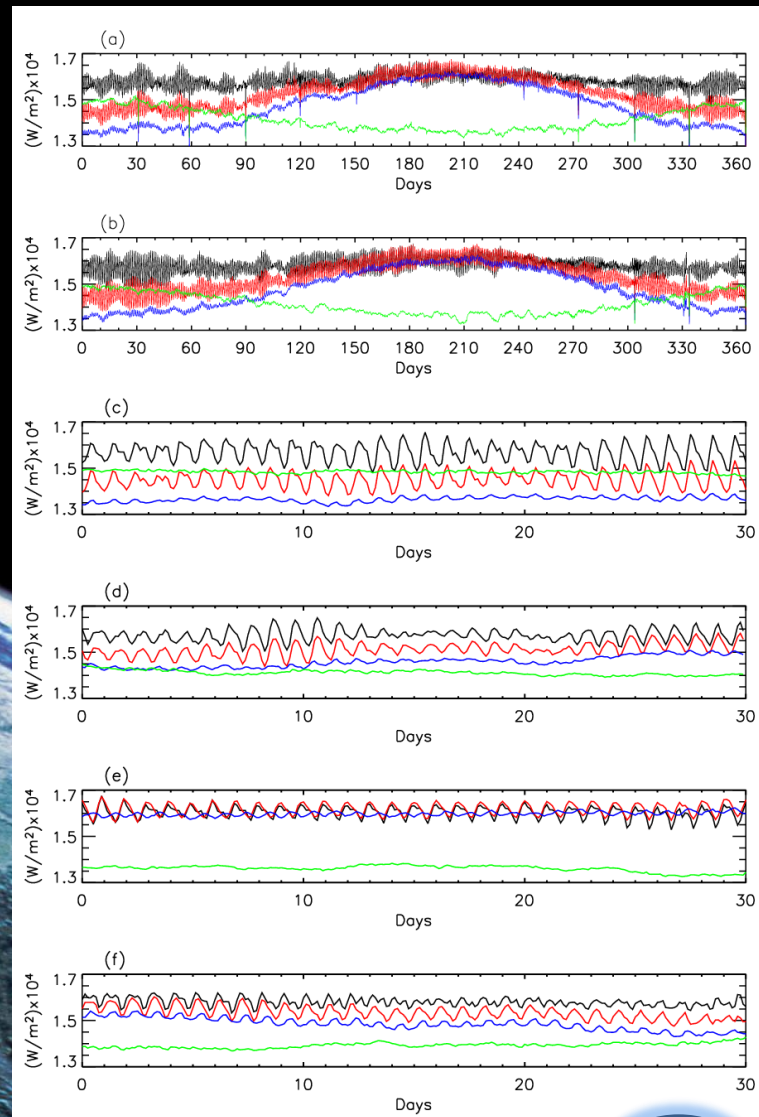
Table 1: Photometric variability of the integrated infrared flux at several time scales and seasons.

Annual Time Series Variability (%)

Inclination angle	1987	2001
0°(Equator)	4	3
45°(Mid-latitude)	12	11
90°(N. Pole)	17	17
-90°(S. Pole)	10	10

Daily Mean(Max)Time Series Variability (%)

Inclination angle	Jan	Apr	Jul	Oct
0°(Equator)	7(11)	4(9)	4(8)	3(6)
45°(Mid-latitude)	7(8)	4(8)	4(8)	4(6)
90°(N. Pole)	1(2)	1(2)	1(2)	2(3)
-90°(S. Pole)	1(2)	0(1)	0(2)	0(1)



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Light Curves

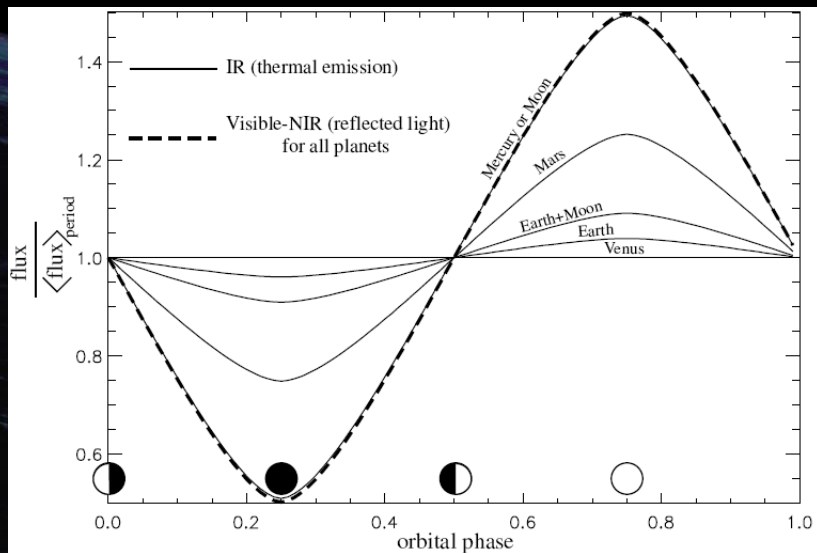
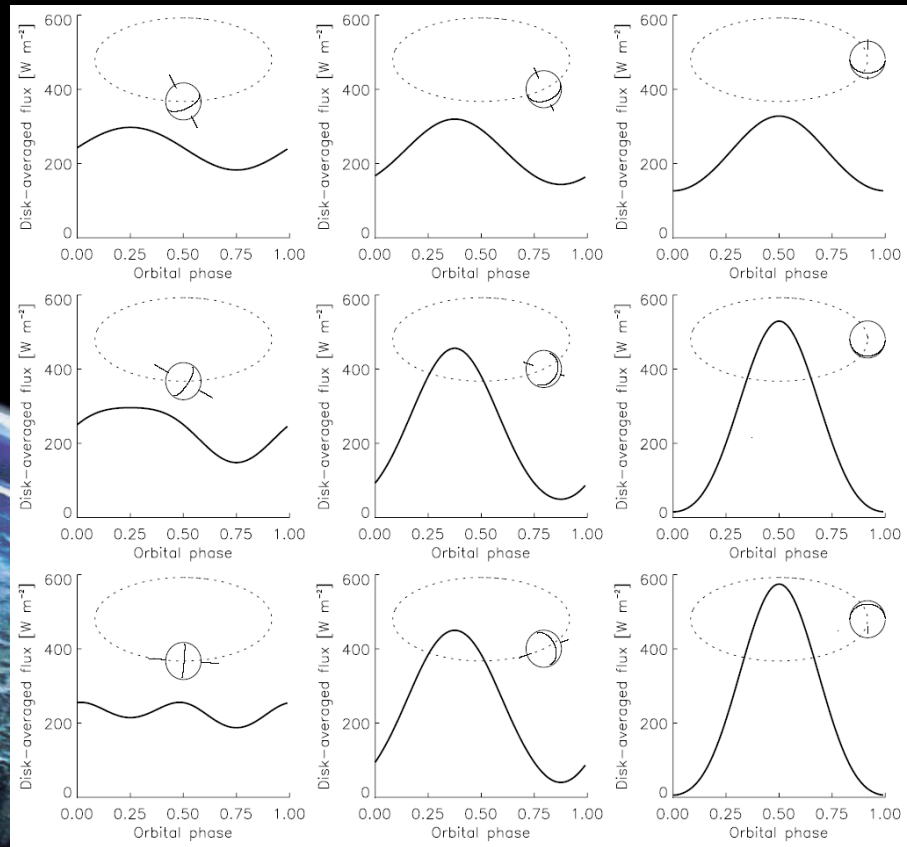


Figure 2. Variation of IR flux over an orbital period for different kinds of planets. The reflected light received from a planet is directly proportional to the illuminated fraction of the planet (visual phase: ϕ). In return, the IR emission only follows the variation of ϕ for atmosphereless planets. In this simple model the planets have null obliquities and circular orbits seen with an inclination of 45° .

Selsis, 2004



Gaidos, 2004

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Light Curves

* Maxima and minima change with season.

* Equator and mid-latitude viewing angles show variation with day.

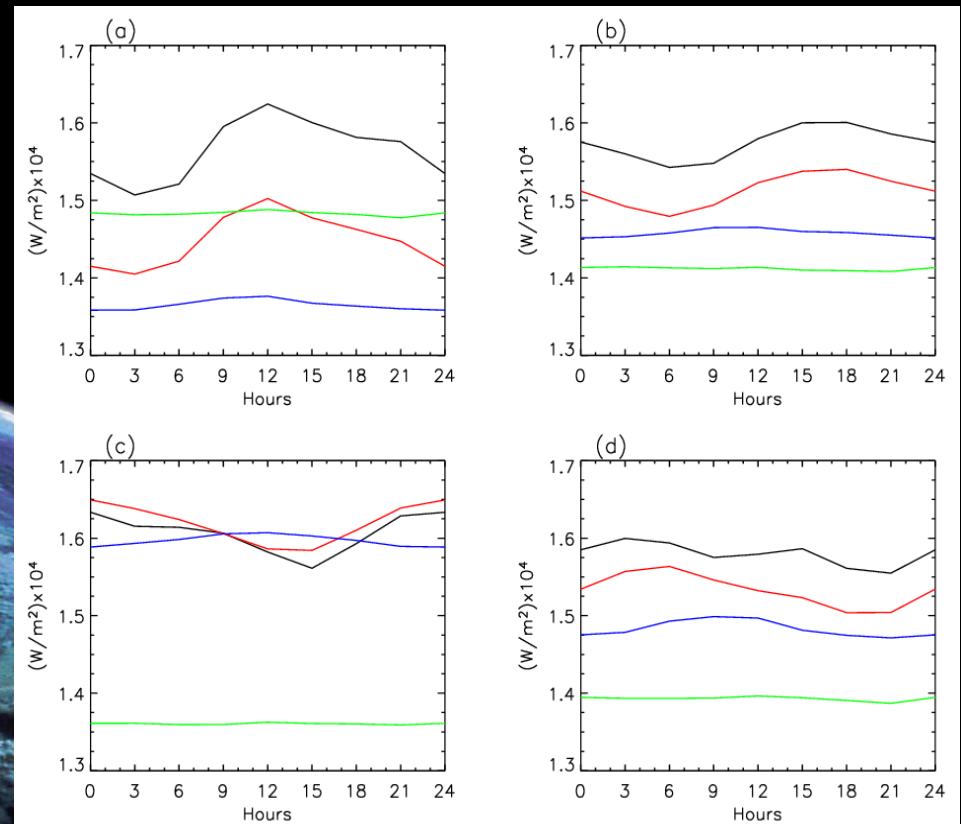


Fig. 3.— Light curves. From upper left to lower right corner: January, April, July and October 2001 for 90°, 45°, 0° and -90° inclination angles (black, red, blue and green lines respectively).

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Autocorrelation

* Rotation period: 24 hours

* January and July have are more stable.
(Polar fronts, anticyclons)

* April and October are more cloudy
(hurricanes, monsoons)

* Polar regions are more homogeneous.

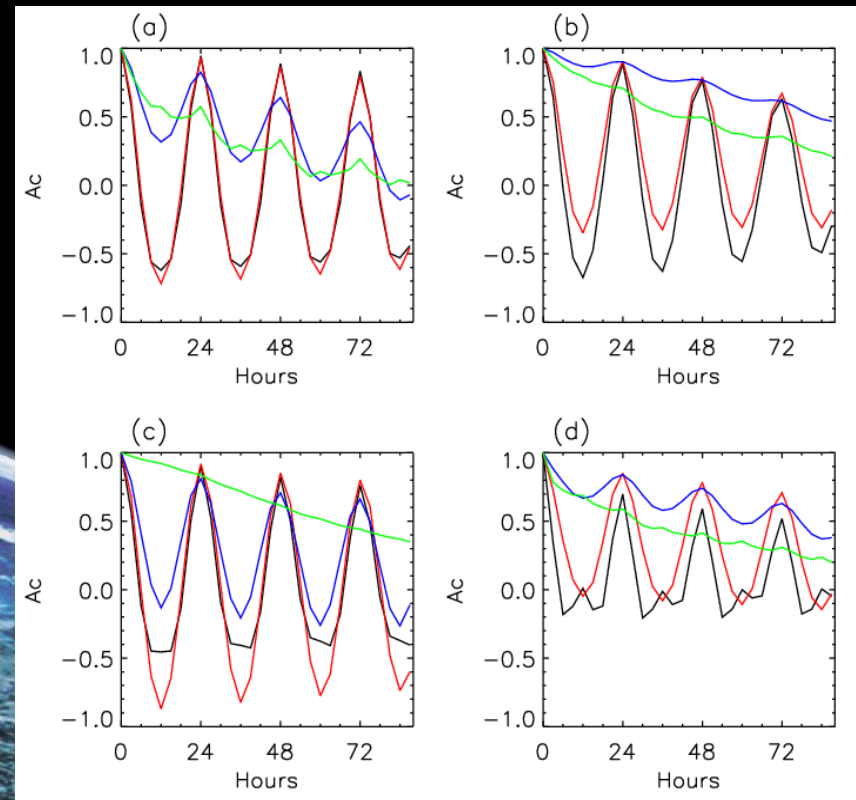
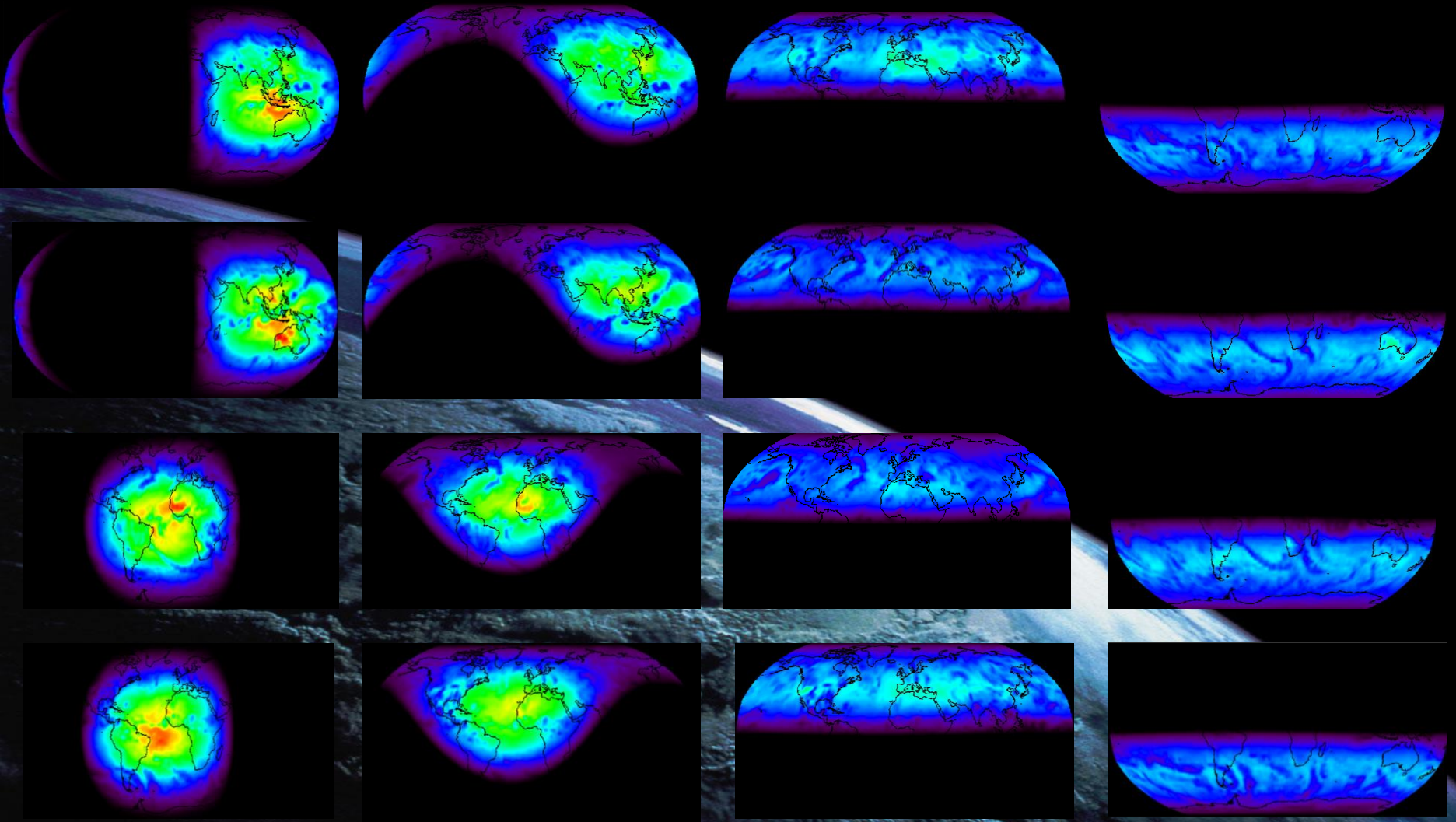


Fig. 4.— Autocorrelation. From upper left to lower right corner: January, April, July and October 2001 for 90°, 45°, 0° and -90° inclination angles (black, red, blue and green lines respectively).

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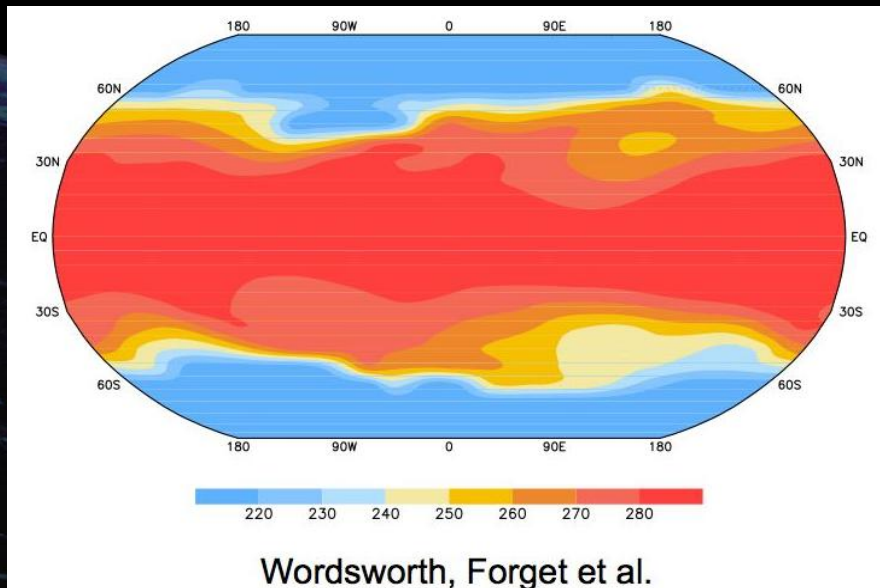
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Conclusions

- * Annual maximum.- Northern hemisphere summer (big continental areas)
- * Daily maximum.- Big continental regions (Sahara desert)
- * Daily minimum.- Oceanic zones (Pacific Ocean)
- * 24 hrs rotation period
- * Cloud lifetime of few days

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Future work ...



Gliese 581d: Carte de température

- 5 bars CO₂
- étoile M ≠ Soleil
- rotation synchrone (toujours la même face éclairée)
- obliquité = 0°

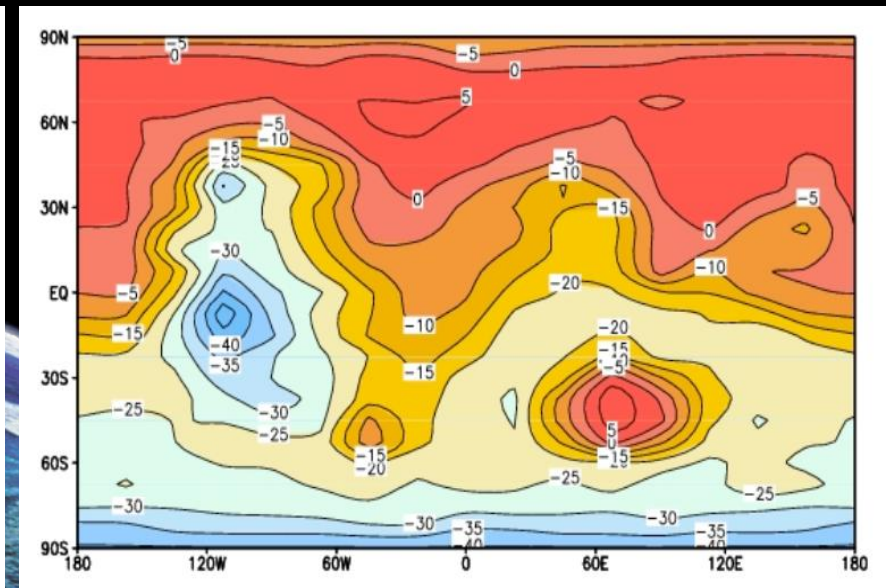
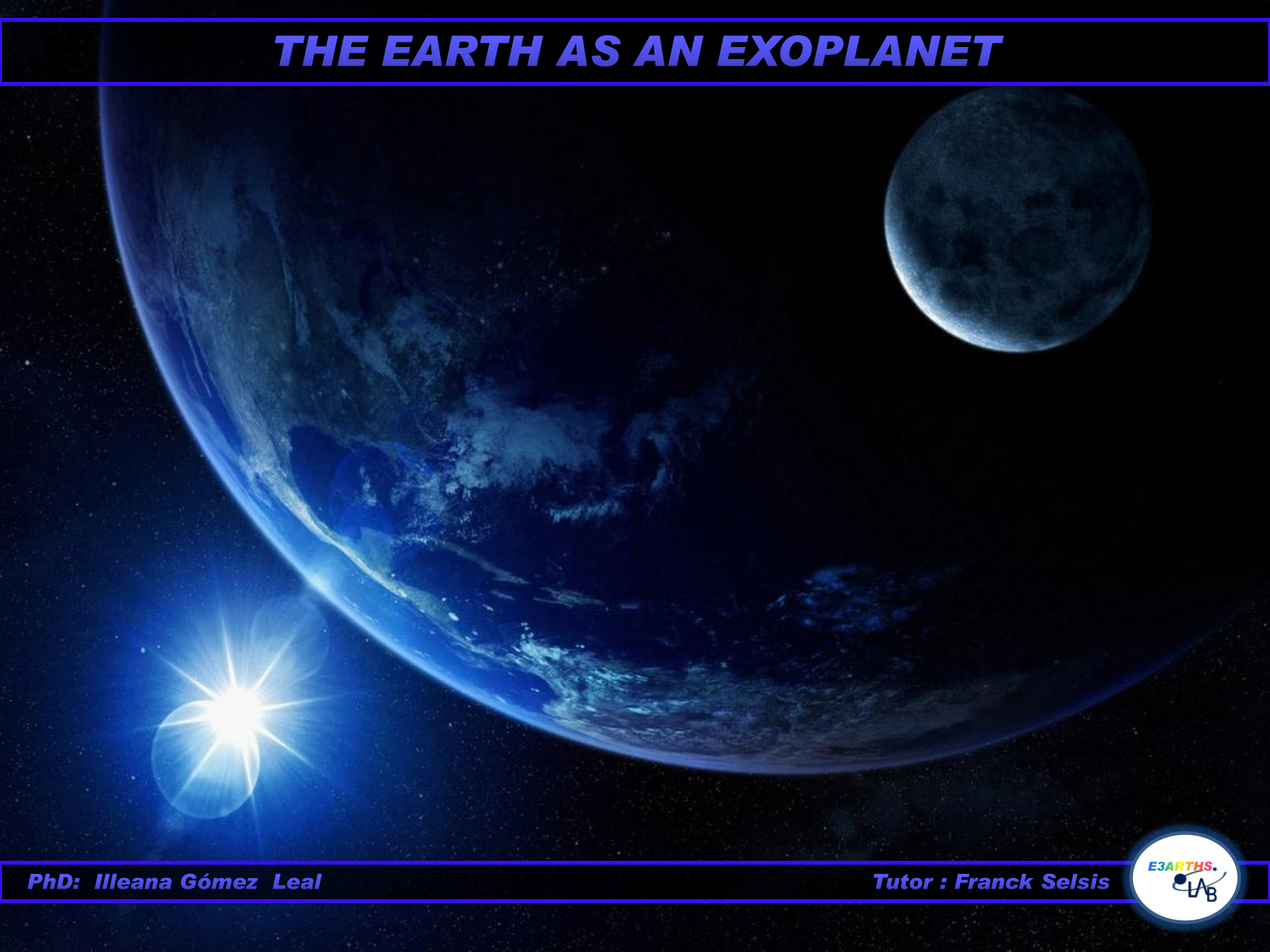


Fig. 3 GCM simulation of early Mars. This temperature map is a result of a Martian simulation at 4 Ga, with a 2 bar CO₂ atmosphere, including CO₂ ice clouds. F. Forget, unpublished.

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